

In the format provided by the authors and unedited.

# A Hubble constant measurement from superluminal motion of the jet in GW170817

K. Hotokezaka<sup>1\*</sup>, E. Nakar<sup>2\*</sup>, O. Gottlieb<sup>2</sup>, S. Nissanke<sup>3,4,5</sup>, K. Masuda<sup>1</sup>, G. Hallinan<sup>6</sup>,  
K. P. Mooley<sup>6,7</sup> and A. T. Deller<sup>8,9</sup>

<sup>1</sup>Department of Astrophysical Sciences, Princeton University, Princeton, NJ, USA. <sup>2</sup>The Raymond and Beverly Sackler School of Physics and Astronomy, Tel Aviv University, Tel Aviv, Israel. <sup>3</sup>GRAPPA, Anton Pannekoek Institute for Astronomy and Institute of High-Energy Physics, University of Amsterdam, Amsterdam, The Netherlands. <sup>4</sup>Nikhef, Amsterdam, The Netherlands. <sup>5</sup>Department of Astrophysics/IMAPP, Radboud University Nijmegen, Nijmegen, The Netherlands. <sup>6</sup>Caltech, Pasadena, CA, USA. <sup>7</sup>National Radio Astronomy Observatory, Socorro, NM, USA. <sup>8</sup>Centre for Astrophysics and Supercomputing, Swinburne University of Technology, Hawthorn, Victoria, Australia. <sup>9</sup>ARC Centre of Excellence for Gravitational Wave Discovery (OzGrav), Hawthorn, Victoria, Australia. \*e-mail: [kentah@astro.princeton.edu](mailto:kentah@astro.princeton.edu); [udini@wise.tau.ac.il](mailto:udini@wise.tau.ac.il)

# A Hubble constant measurement from superluminal motion of the jet in GW170817

K. Hotokezaka<sup>1</sup>, E. Nakar<sup>2</sup>, O. Gottlieb<sup>2</sup>, S. Nissanke<sup>3,4,5</sup>, K. Masuda<sup>1,10</sup>, G. Hallinan<sup>6</sup>, K. P. Mooley<sup>6,7</sup>, & A. T. Deller<sup>8,9</sup>

<sup>1</sup>*Department of Astrophysical Sciences, Princeton University, Peyton Hall, Princeton, NJ 08544, USA*

<sup>2</sup>*The Raymond and Beverly Sackler School of Physics and Astronomy, Tel Aviv University, Tel Aviv 69978, Israel*

<sup>3</sup>*GRAPPA, Anton Pannekoek Institute for Astronomy and Institute of High-Energy Physics, University of Amsterdam, Science Park 904, 1098 XH Amsterdam, The Netherlands*

<sup>4</sup>*Nikhef, Science Park 105, 1098 XG Amsterdam, The Netherlands*

<sup>5</sup>*Department of Astrophysics/IMAPP, Radboud University Nijmegen, P.O. Box 9010, 6500 GL Nijmegen, The Netherlands*

<sup>6</sup>*Caltech, 1200 E California Blvd, MC 249-17, Pasadena, CA 91125, USA*

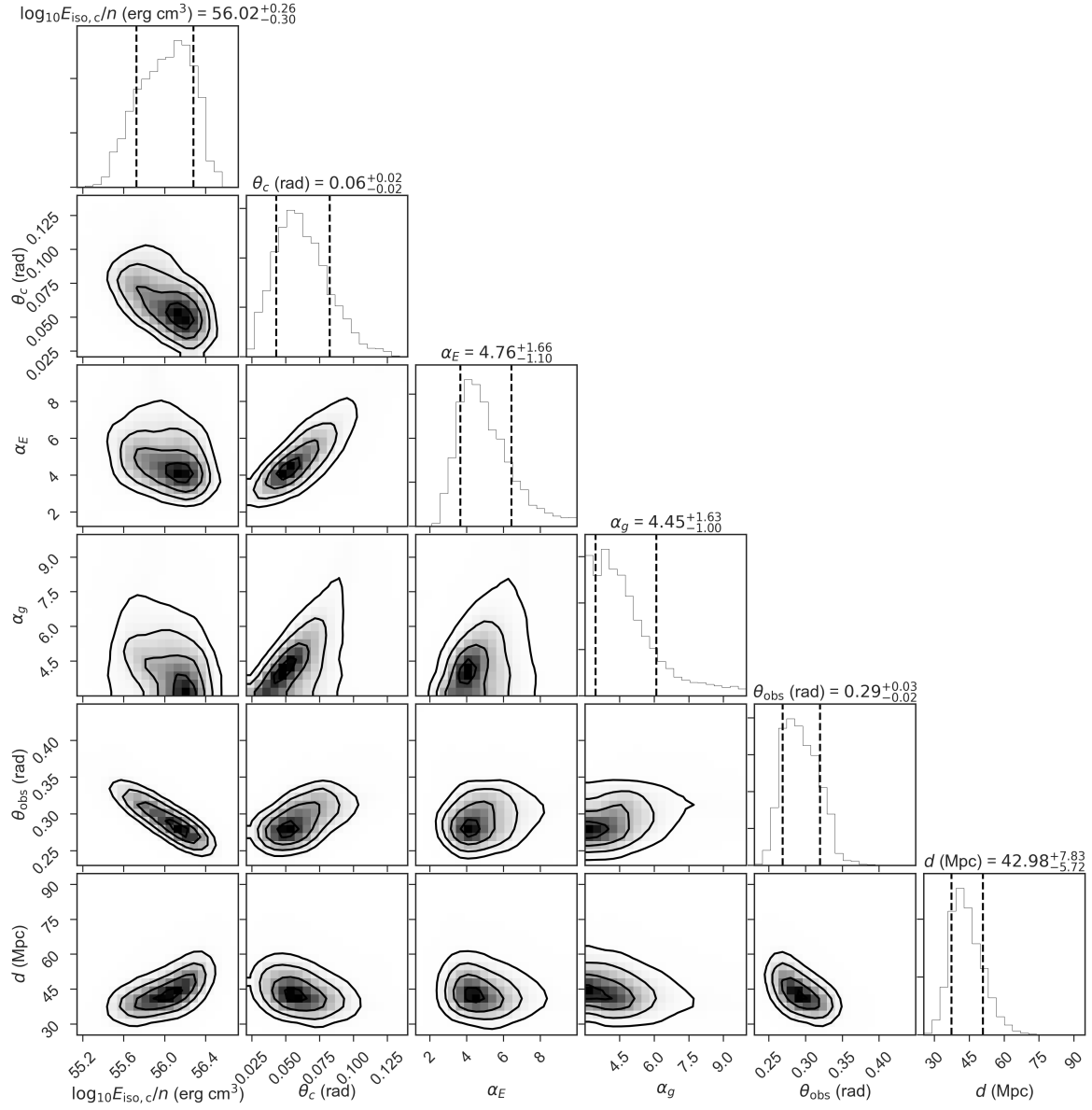
<sup>7</sup>*National Radio Astronomy Observatory, Socorro, New Mexico, 87801, USA*

<sup>8</sup>*Centre for Astrophysics & Supercomputing, Swinburne University of Technology, John St, Hawthorn VIC 3122 Australia*

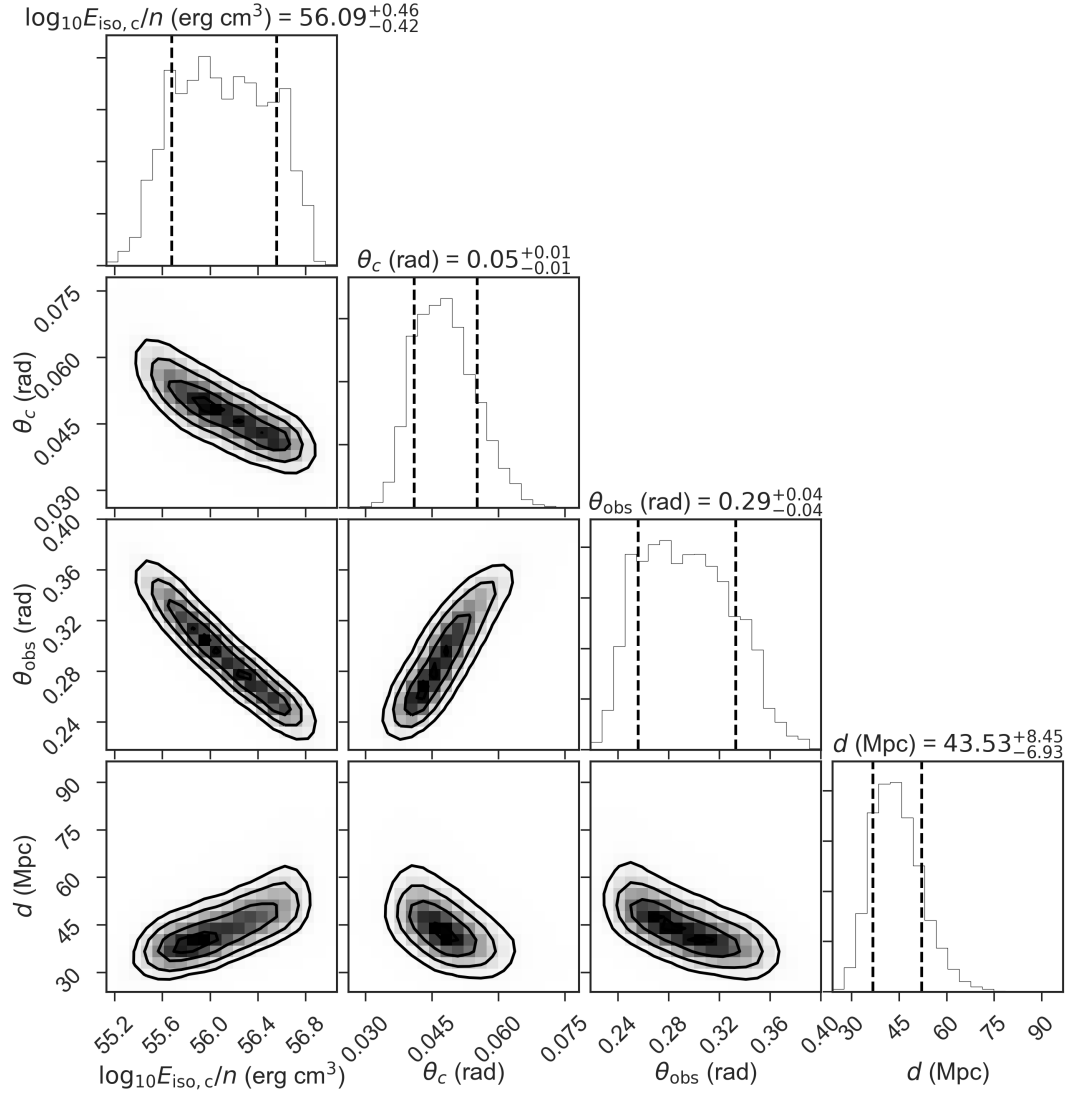
<sup>9</sup>*ARC Centre of Excellence for Gravitational Wave Discovery (OzGrav), Australia*

<sup>10</sup>*NASA Sagan Fellow*

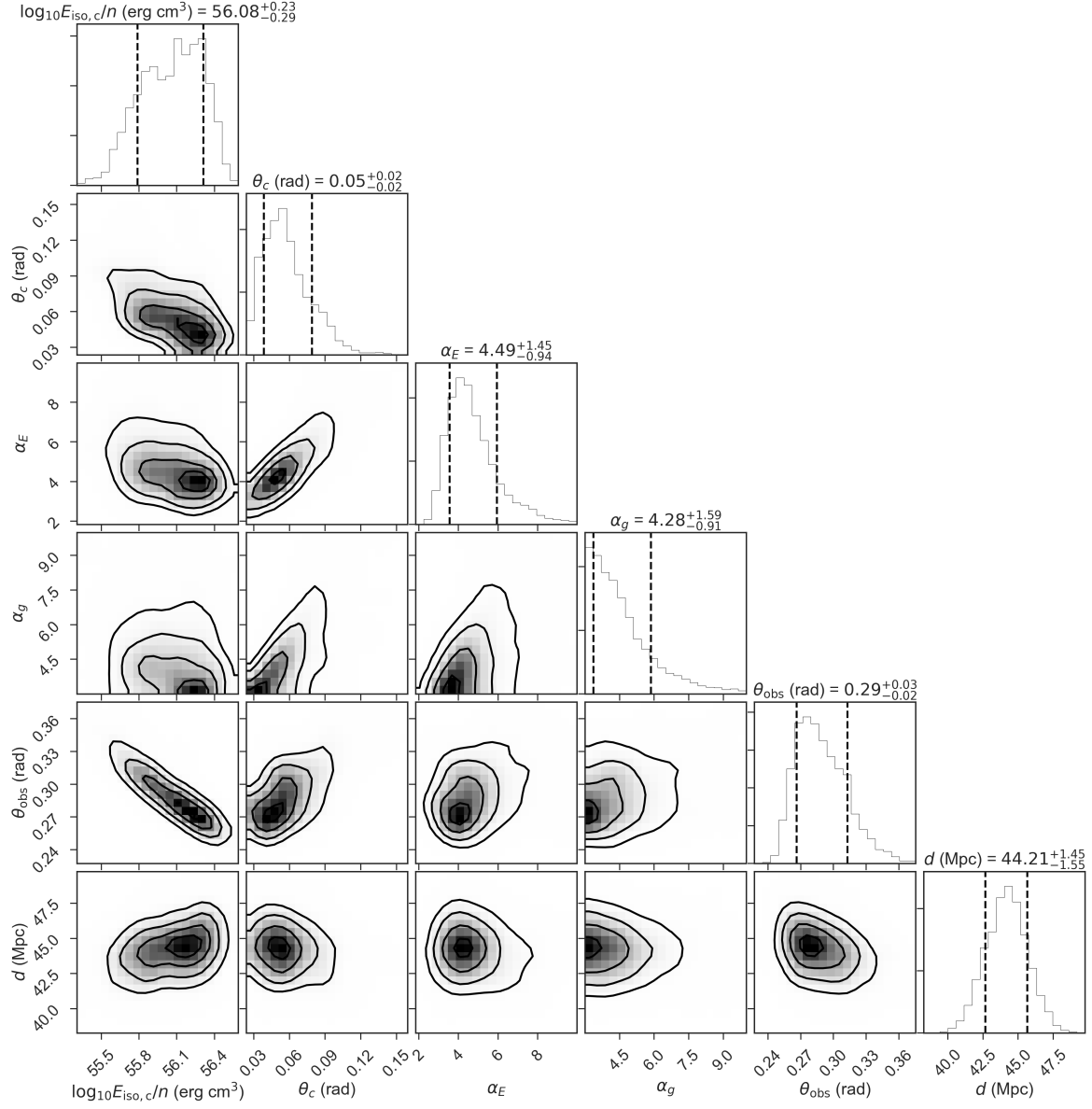
## **Supplementary Information**



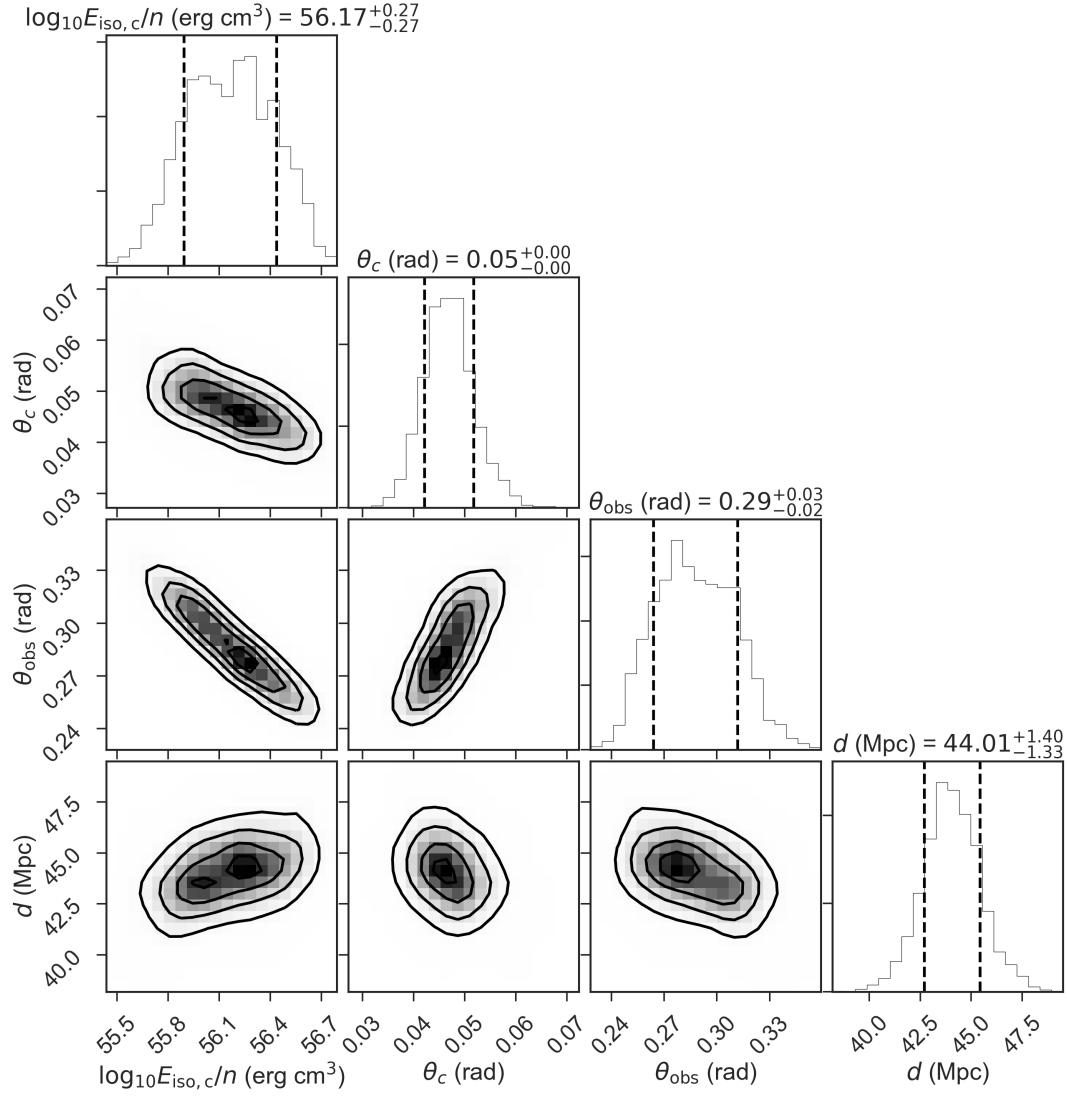
**Supplementary Figure 1: Corner plot<sup>1</sup> for a Power-Law Jet model.** The afterglow light curve at 3 GHz and the centroid motion resolved by VLBI are used as the observed input data. Vertical lines depict 68% credible intervals.



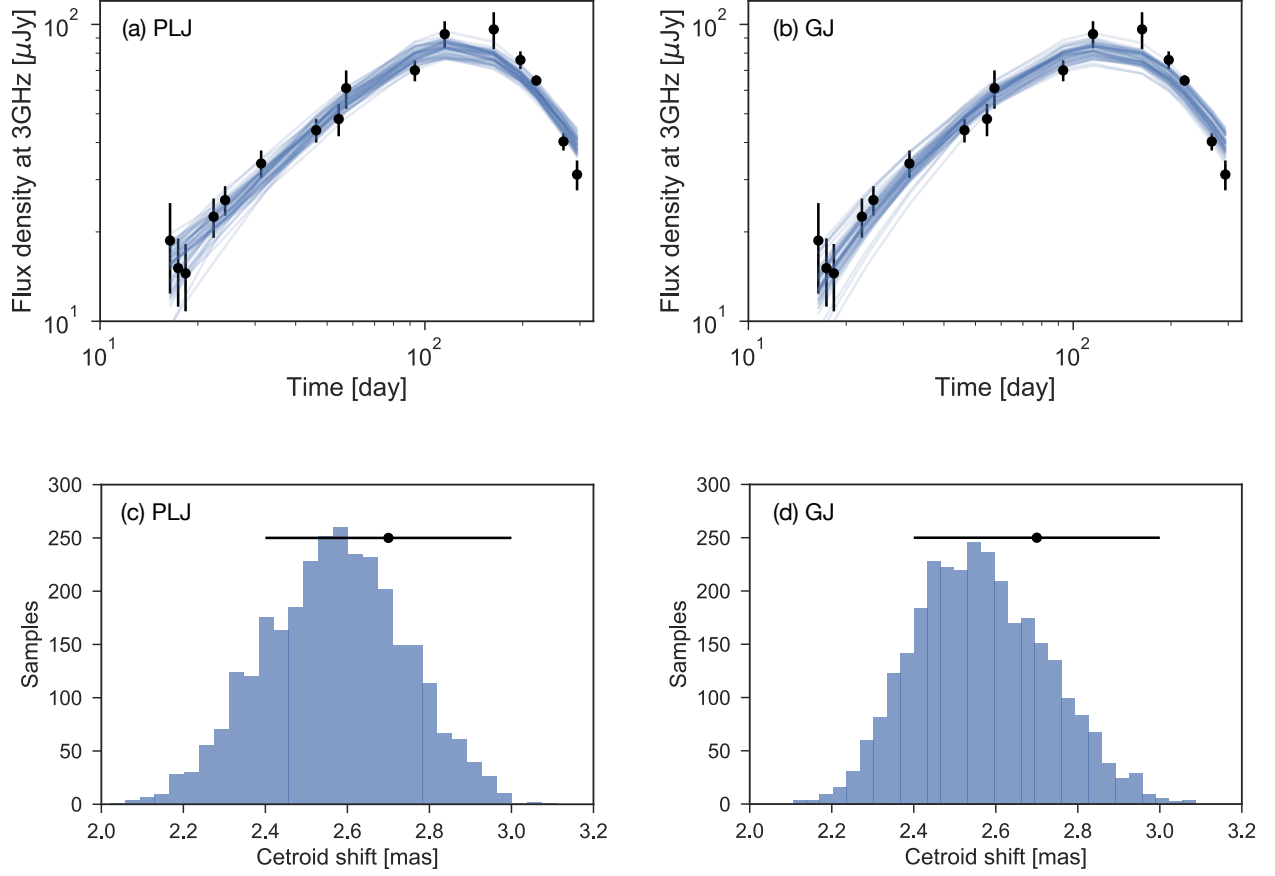
**Supplementary Figure 2: Same as Figure 1 but for a Gaussian Jet model.**



**Supplementary Figure 3: Corner plot for the combined GW-EM analysis with a Power-Law Jet model.** The afterglow light curve at 3 GHz and the centroid motion resolved by VLBI are used as the observed input data. Vertical lines depict 68% credible intervals. Here we use high spin PhenomNR posterior.

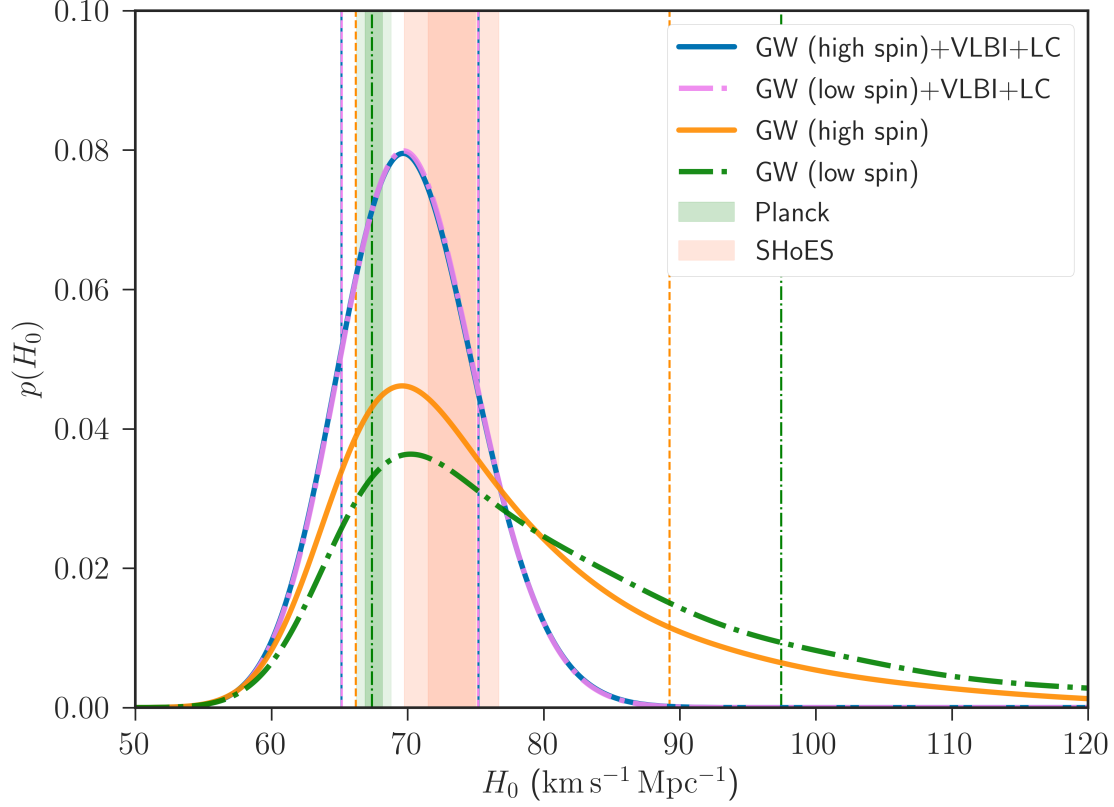


**Supplementary Figure 4: Same as Figure 3 but for a Gaussian Jet model.**



**Supplementary Figure 5: Afterglow light curve at 3 GHz and centroid motion from day 75 to 230 with  $1\sigma$  uncertainties.** Also shown are the light curves calculated with (a) a PLJ and (b) a GJ model, where 50 sets of the model parameters are randomly chosen from the MCMC samples. *Bottom* panels show the histogram of the centroid motion with 3000 samples randomly chosen: (c) a PLJ model and (d) a GJ model. These are the results of the combined GW-VLBI-LC analysis.





**Supplementary Figure 6: Comparison between the  $H_0$  posteriors of the high and low spin priors.** Here we use hydrodynamics simulation jet model ( $0.25 < \theta_{\text{obs}} \left( \frac{d}{41 \text{ Mpc}} \right) < 0.5 \text{ rad}$ ). Solid blue curve: the GW-VLBI-LC analysis with the high-spin prior, dash-dotted purple curve: the GW-VLBI-LC analysis with the low-spin prior, solid orange curve: the GW-only analysis with the high-spin prior, and dash-dotted green curve: the GW-only analysis with the low-spin prior. The vertical lines show symmetric 68% credible interval for each model.

## References

1. Foreman-Mackey, D. corner.py: Scatterplot matrices in python. The Journal of Open Source Software **24** (2016). URL <http://dx.doi.org/10.5281/zenodo.45906>.